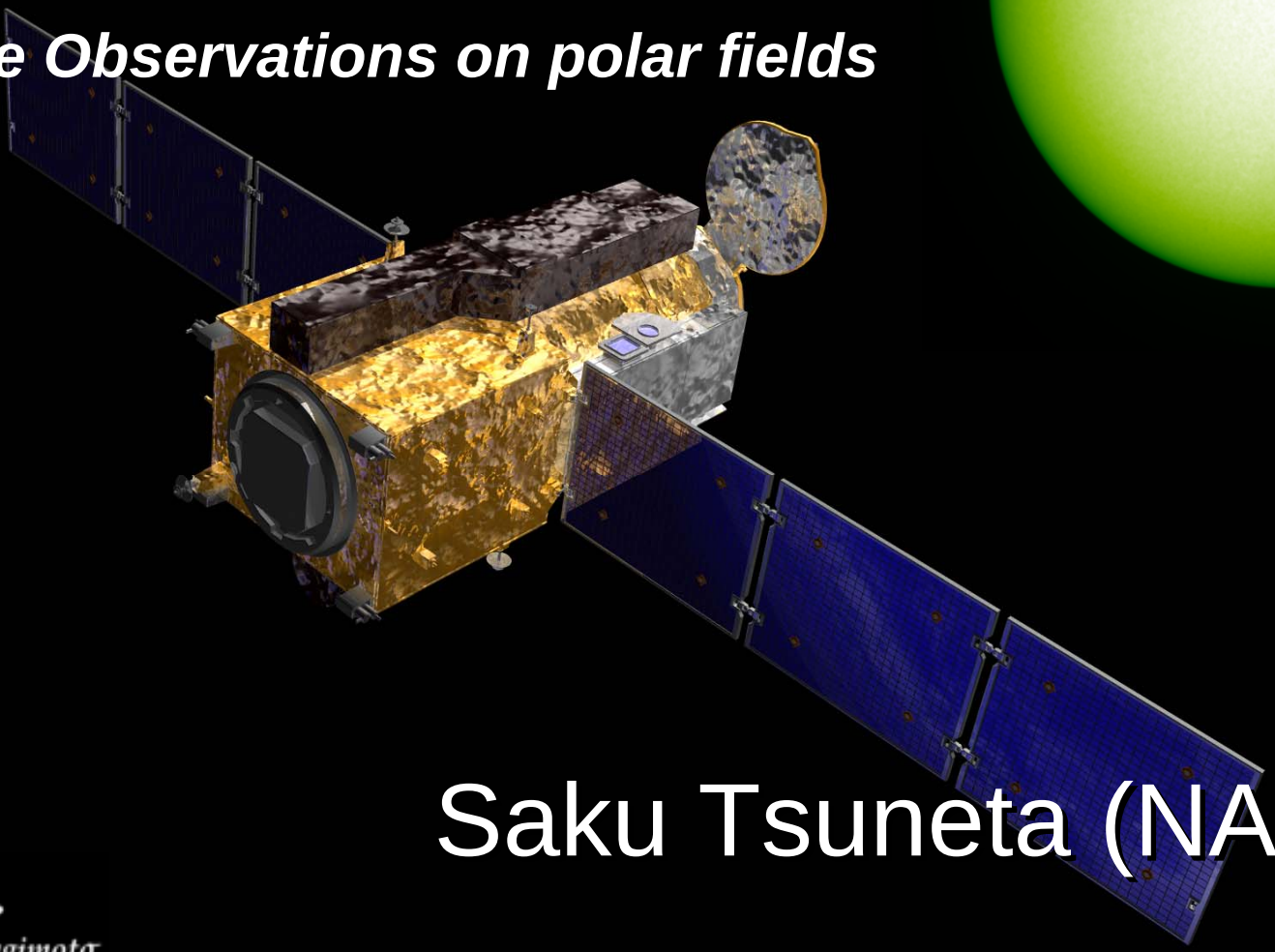
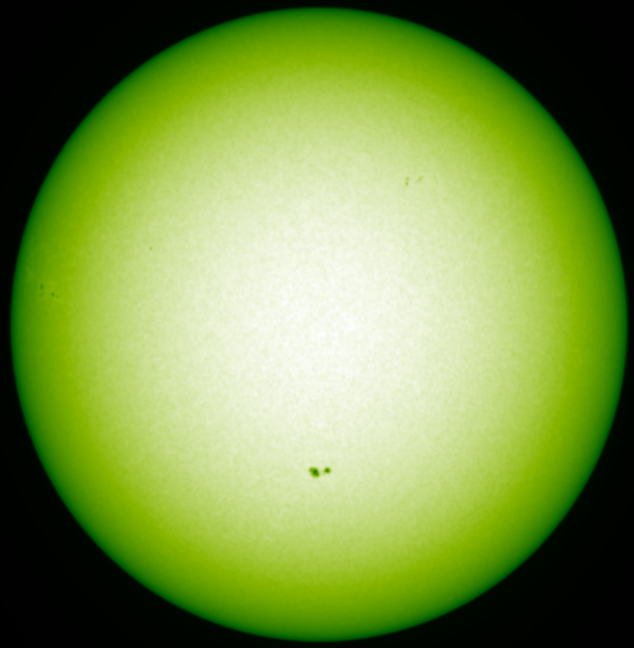
Is the polar region different from the quiet sun?

Hinode Observations on polar fields





Saku Tsuneta (NAOJ)

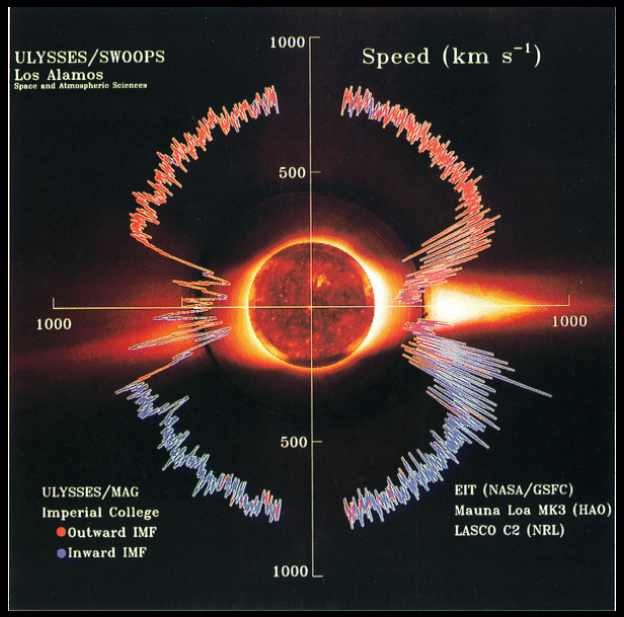
Maki Sugimoto

What is going on in polar region ?

Source of fast solar wind
Location of global poloidal fields
sink of meridional flow

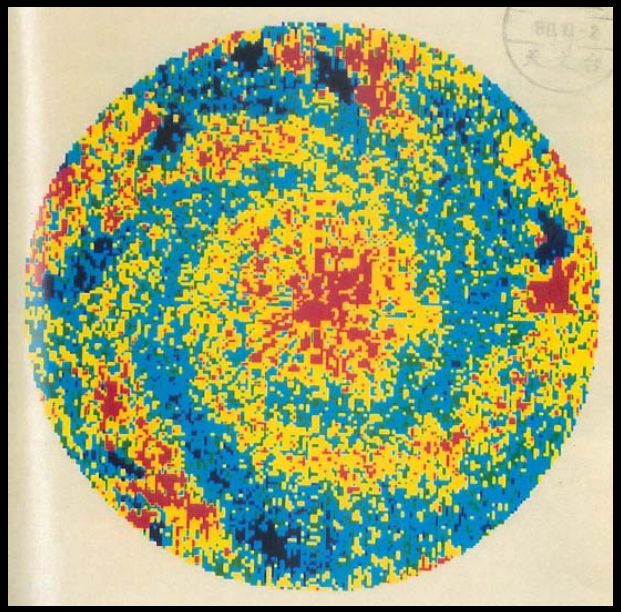
High speed solar wind

Ulysses
(McComas etal 2000)

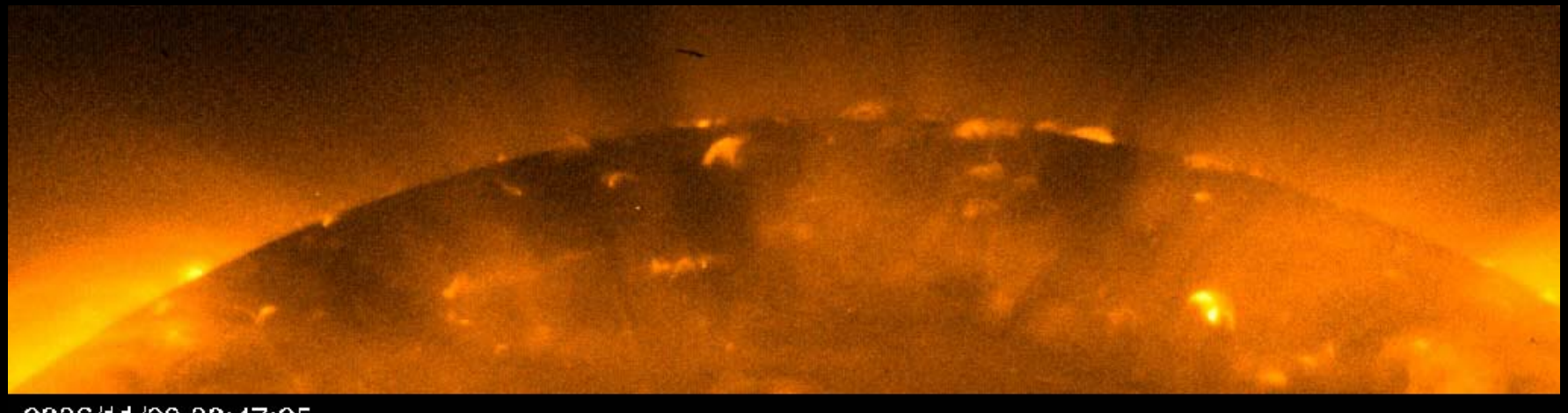


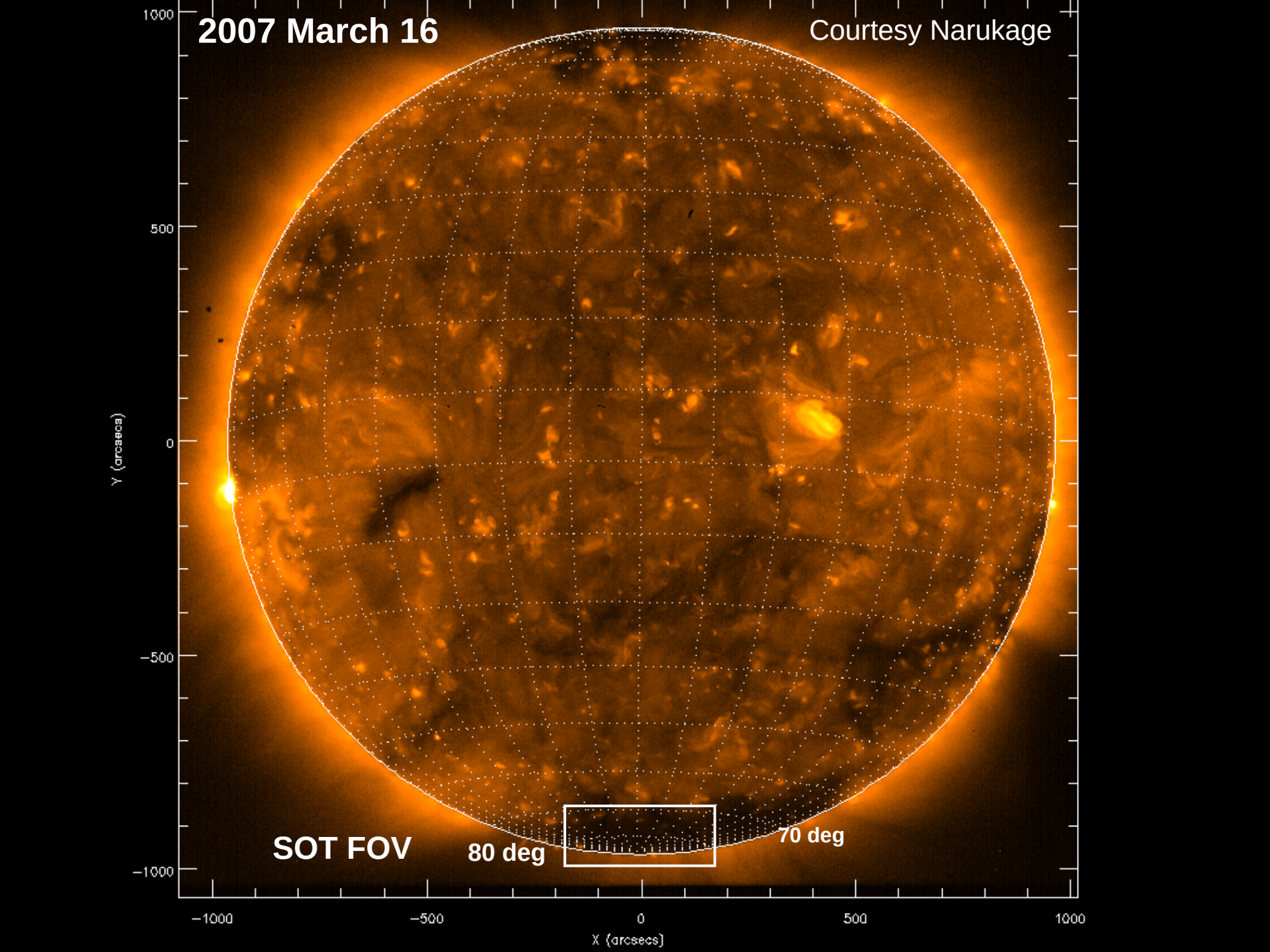
Kitt Peak
(Wang et al)

Polar fields



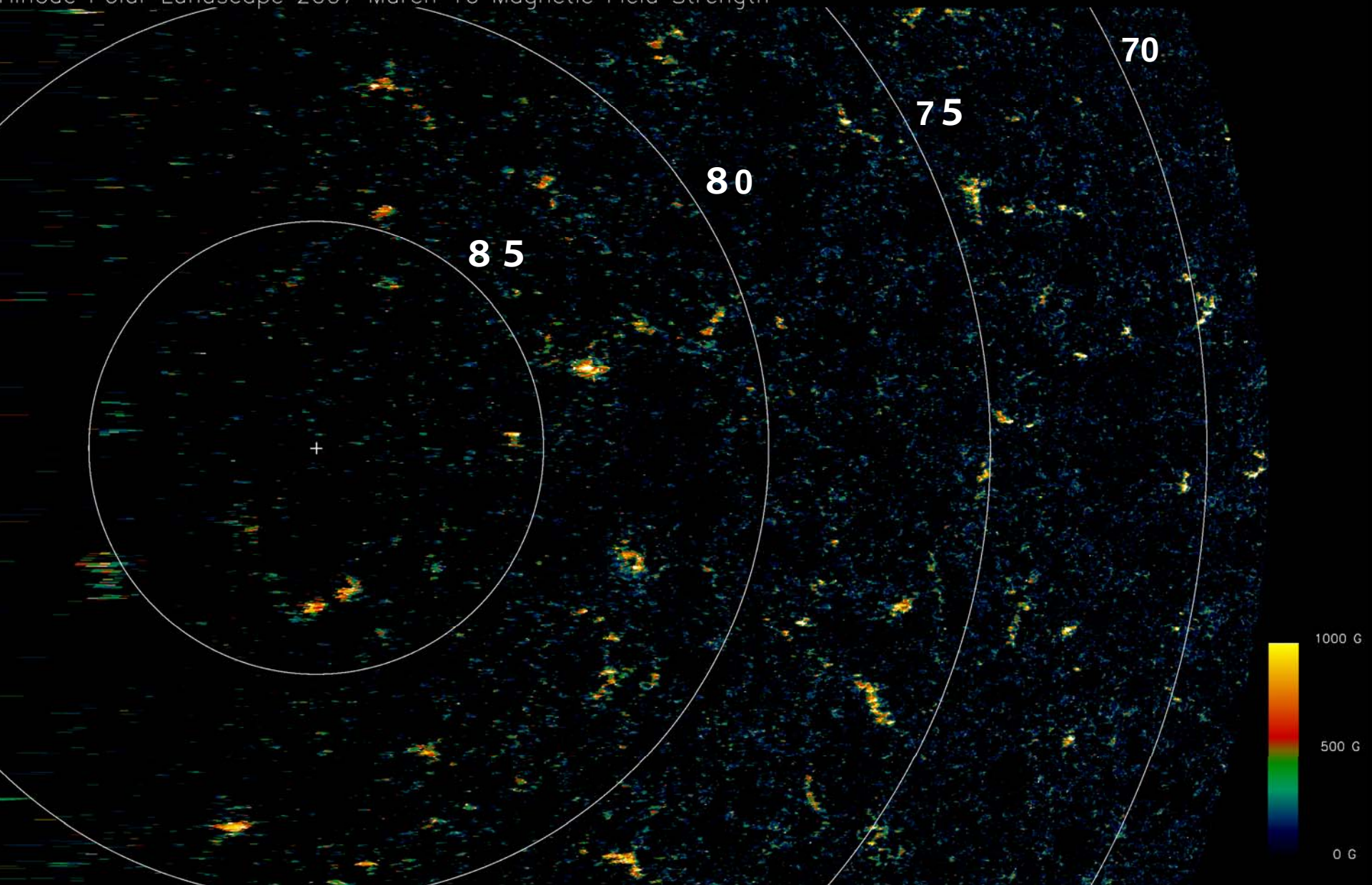
Cirtain etal

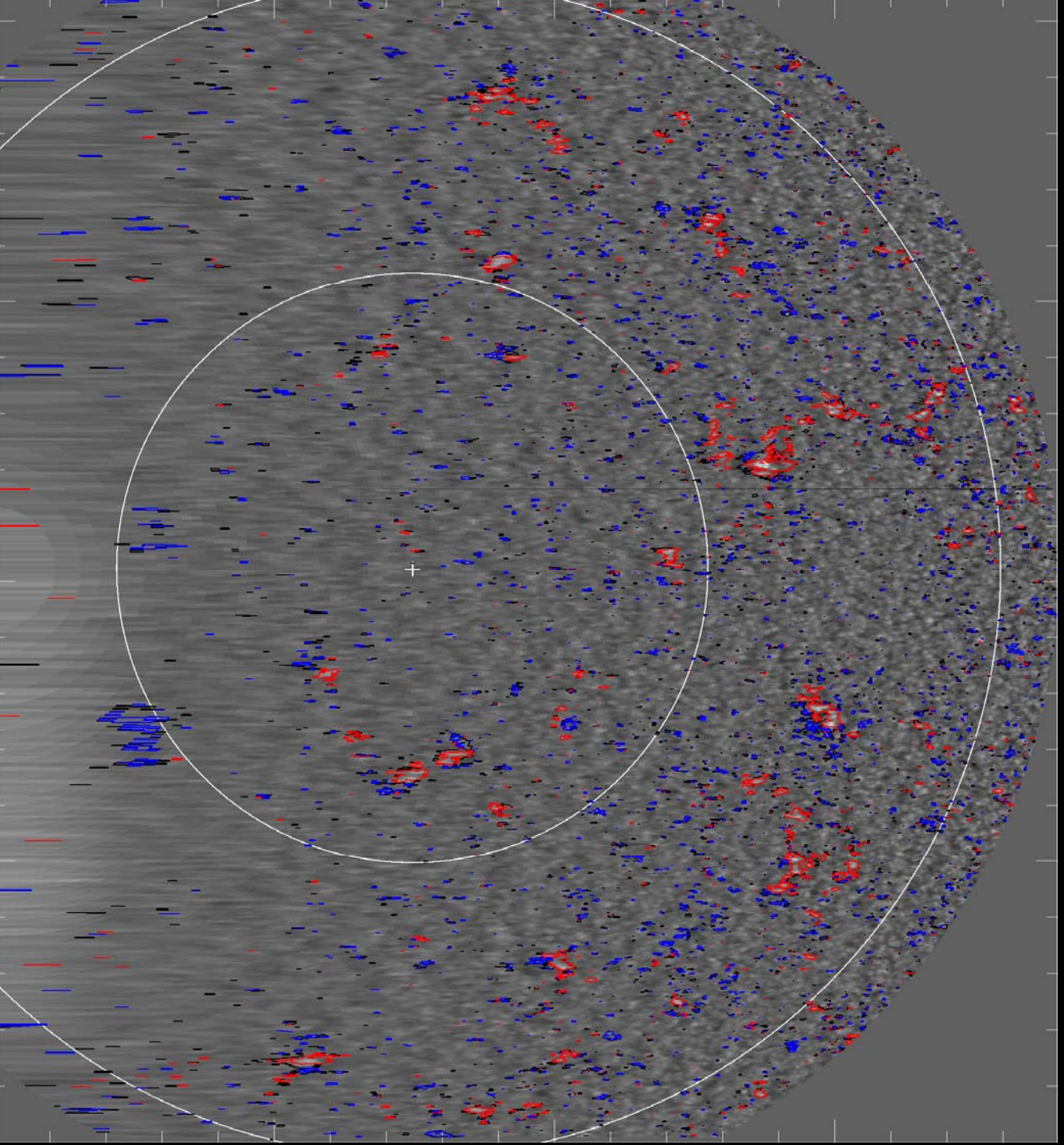



2007 March 16
Courtesy Narukage
SOT FOV
80 deg
70 deg
1000
500
0
−500
−1000
Y (arcsecs)
−1000
−500
0
500
1000
X (arcsecs)

Polar landscape kG field

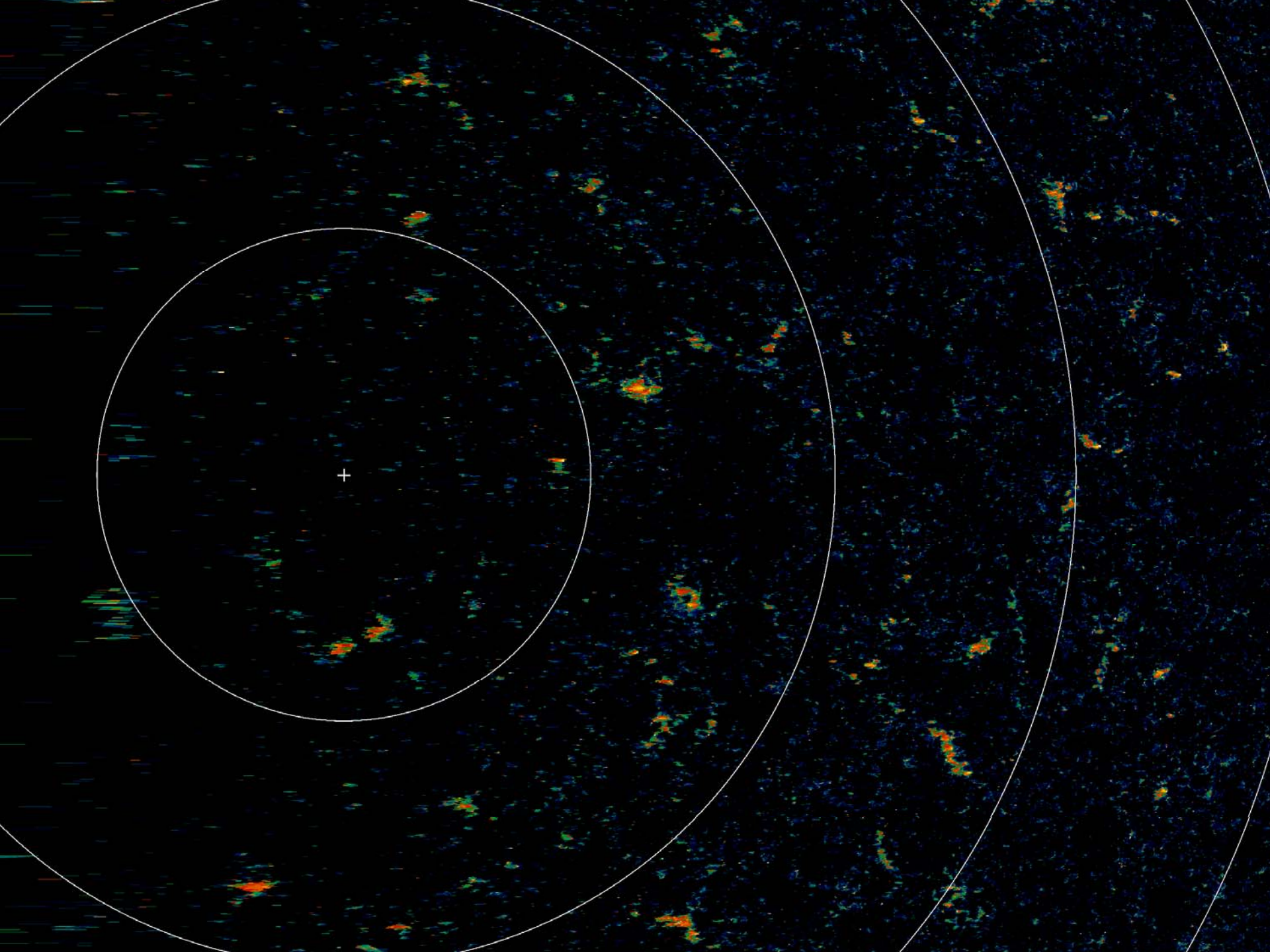
Tsuneta etal (2008 in press)



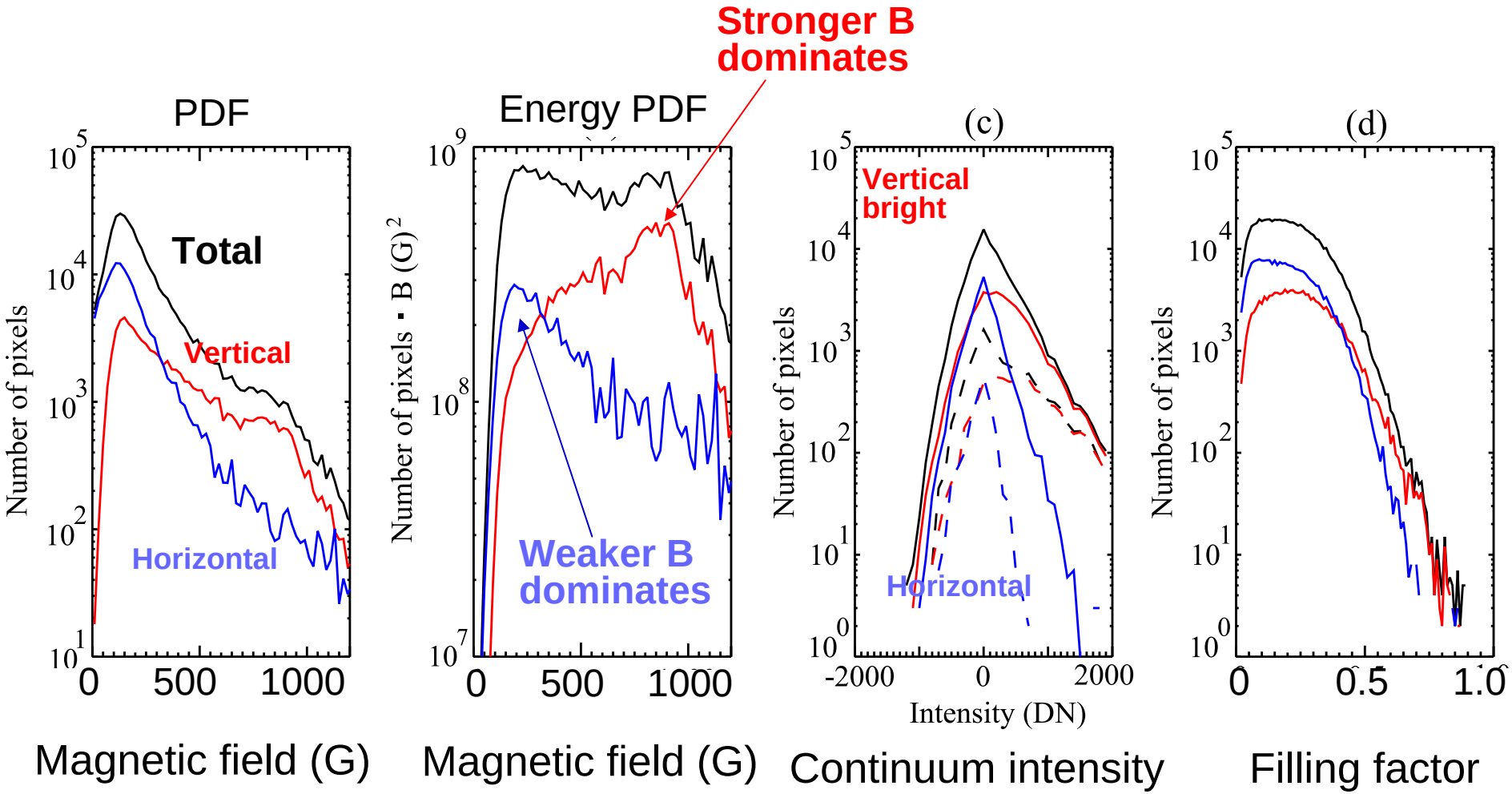


Red：vertical
Blue：horizontal

Polar region consists of same-sign vertical B with horizontal B

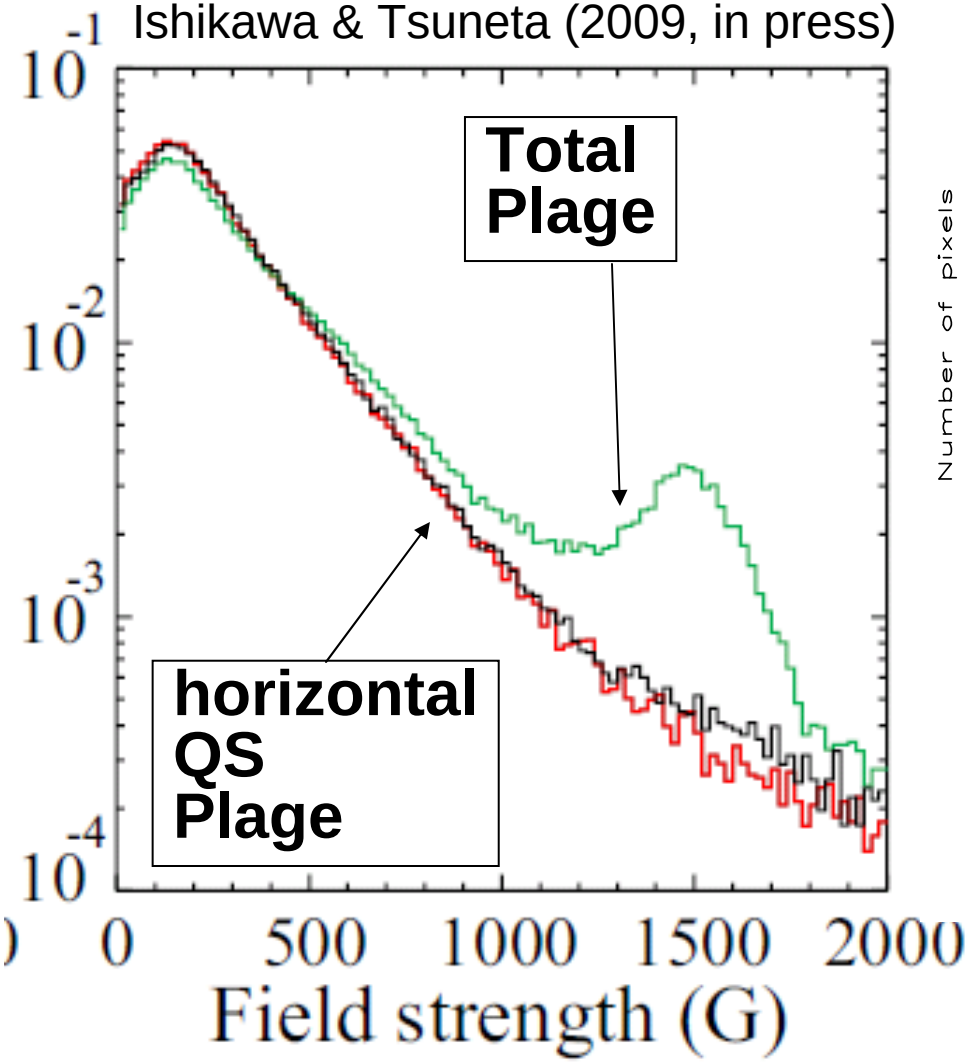


Probability Distribution Function of polar field

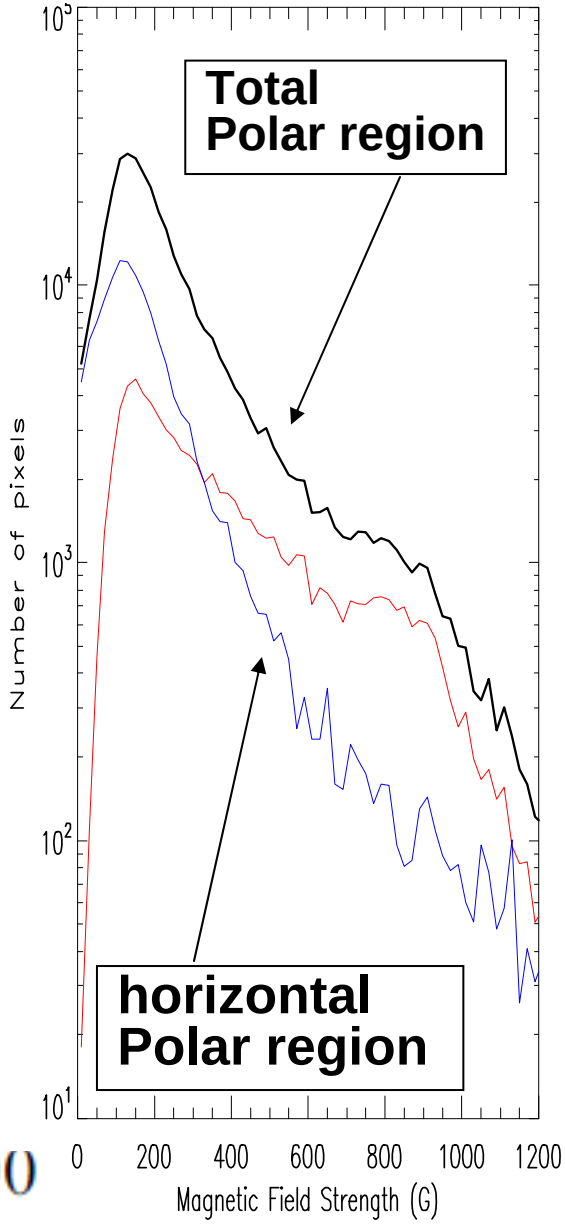


Horizontal field
PDF comparison with Quiet Sun and Plage region

Ishikawa & Tsuneta (2009, in press)



Tsuneta etal (2008 in press)

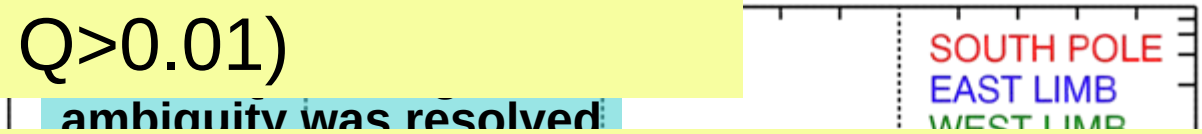


ce between olar region?

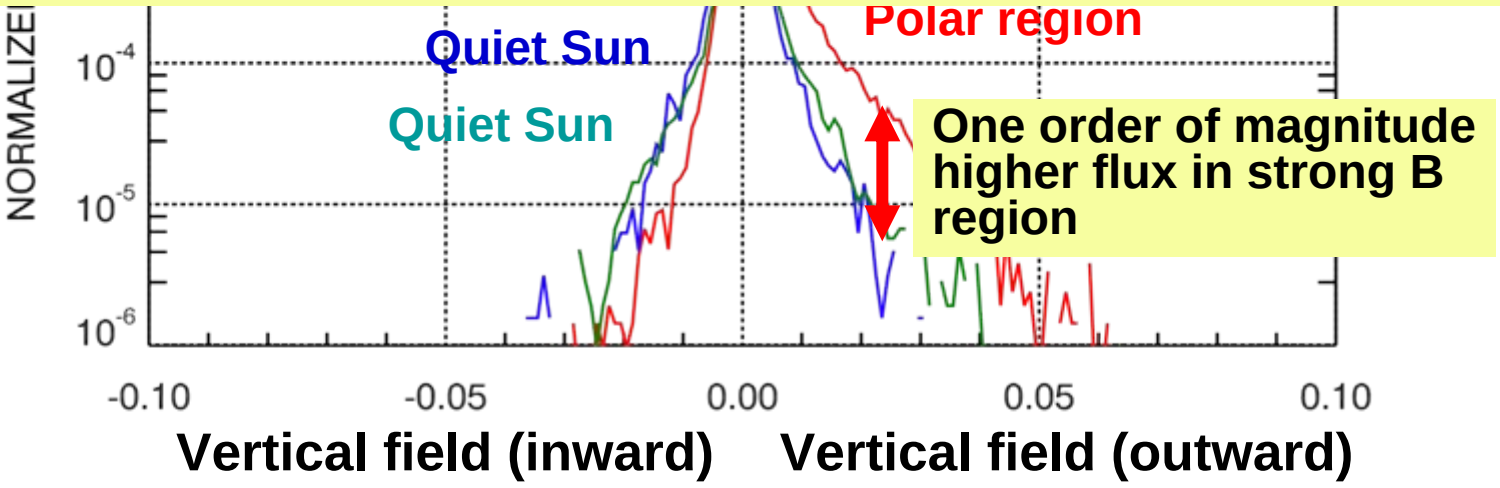
Outward and inward flux
South pole: 96.3% : 3.7%
East limb QS: 50.5% : 49.5%
West limb QS: 64.0% : 36.0%
(Stokes Q>0.01)

Ito & Tsuneta 2008

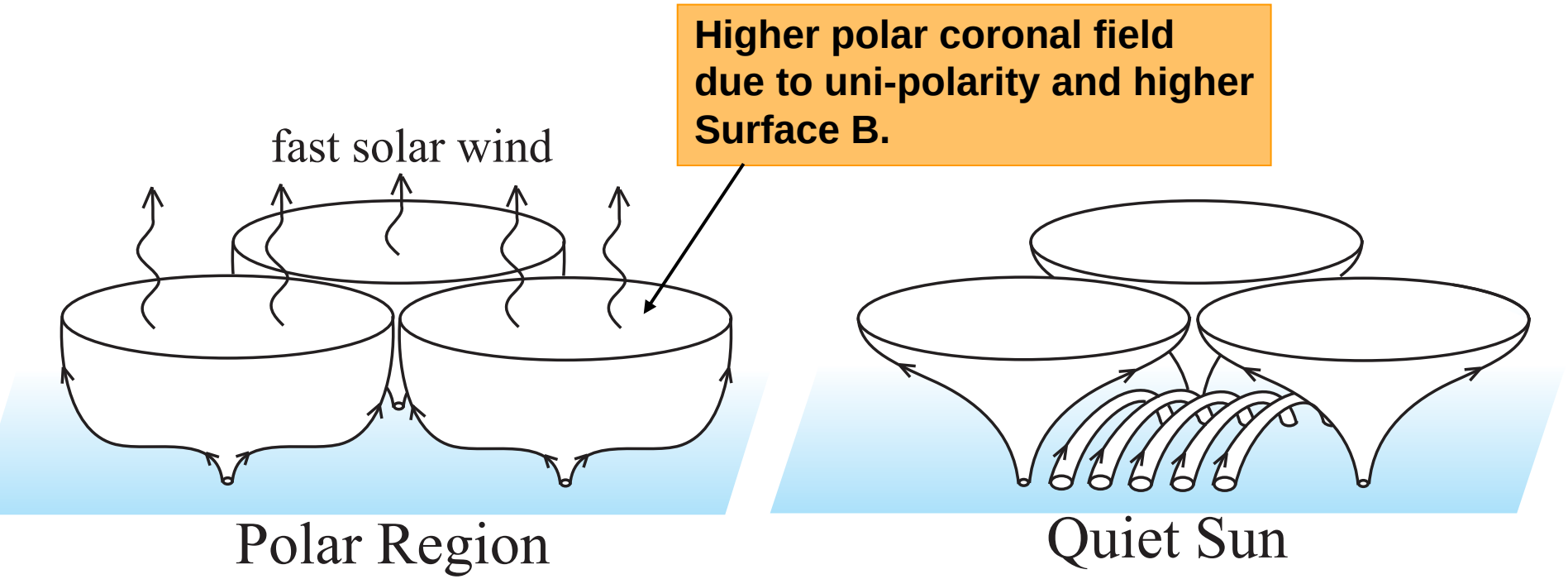
ertical field



- Quiet sun has bipolar field, while polar region is uni-polar.
- Polar region has stronger fields.
- No difference in terms of horizontal fields



Inflation of polar flux tubes near photosphere-chromosphere boundary



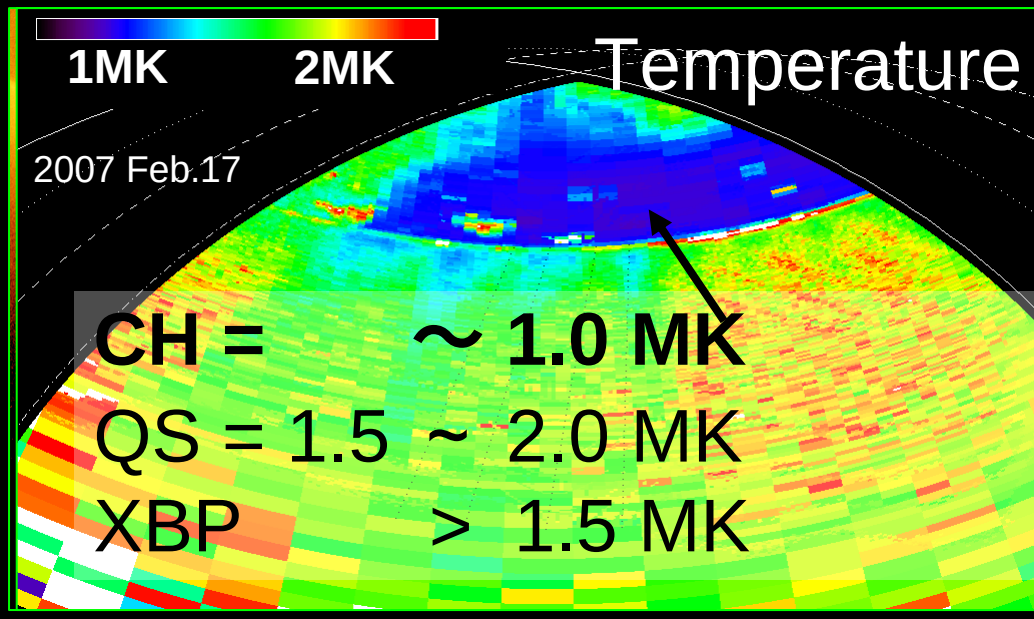
All the flux tubes are connected to interplanetary space.

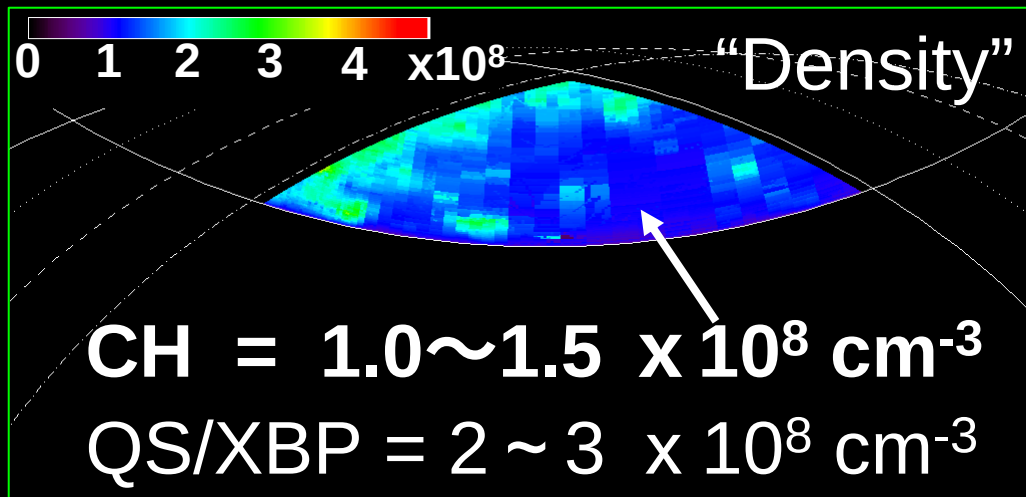
Very small fraction of flux tubes is connected to interplanetary space.

Figure Courtesy Joten Okamoto

Low pressure coronal hole is sustained by higher magnetic pressure of coronal hole

- CH pressure is a factor of 2 (T-diff.) x 3 (n-diff.) = 6 smaller than that of surrounding quiet Sun.
- This must be balanced by stronger CH magnetic field.



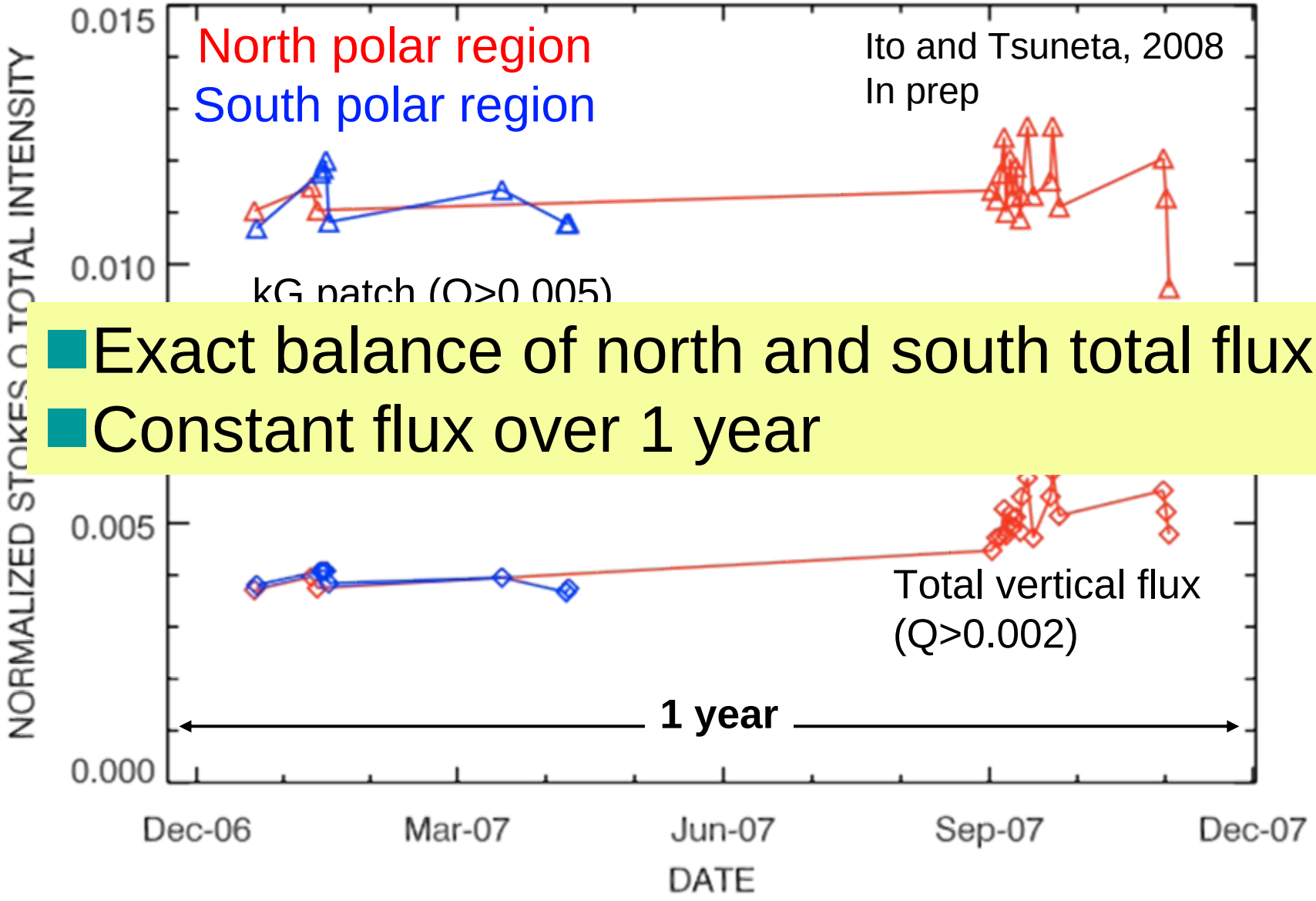


Kano et al (2008)

Total and average magnetic flux

- Magnetic filling factor: **$0.05 < f < 0.35$**
- Total magnetic flux in a pixel
  - *B x f x pixel size with foreshortening correction*
- Total magnetic flux in SOT FOV
  - vertical flux **$2.2\text{-}9.9 \times 10^{21}$ Mx**
  - horizontal flux **$0.4\text{-}2.0 \times 10^{22}$ Mx**
- *Horizontal flux factor of 2 larger*
- Average vertical flux 10.0 G
- Total vertical magnetic flux for > 70 degree
  - **$0.6\text{-}2.5 \times 10^{22}$ Mx**
- Ulysses observations
  - **2×10^{22} Mx** (above 35 degree)

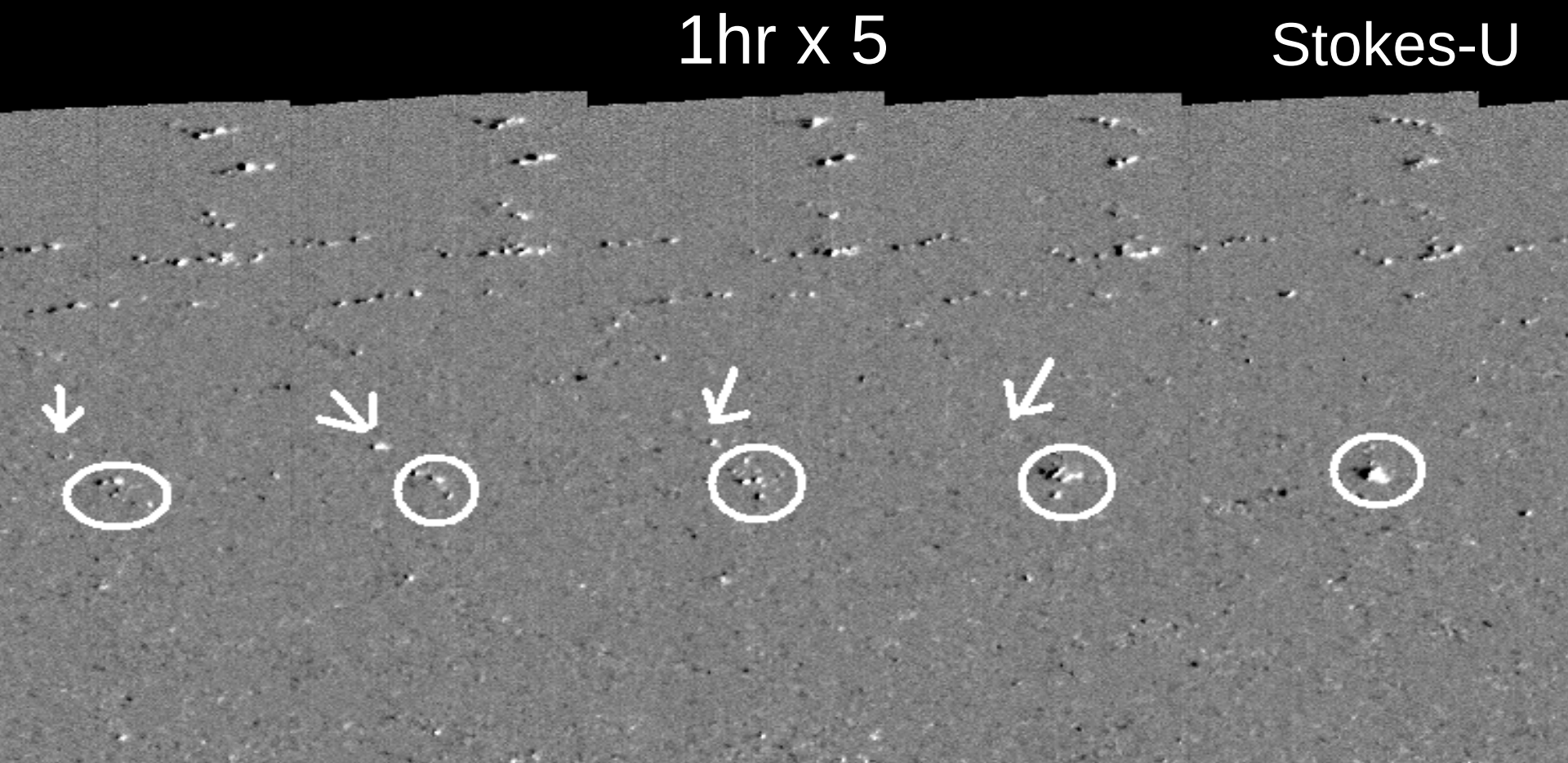
Variation of total polar flux over one year



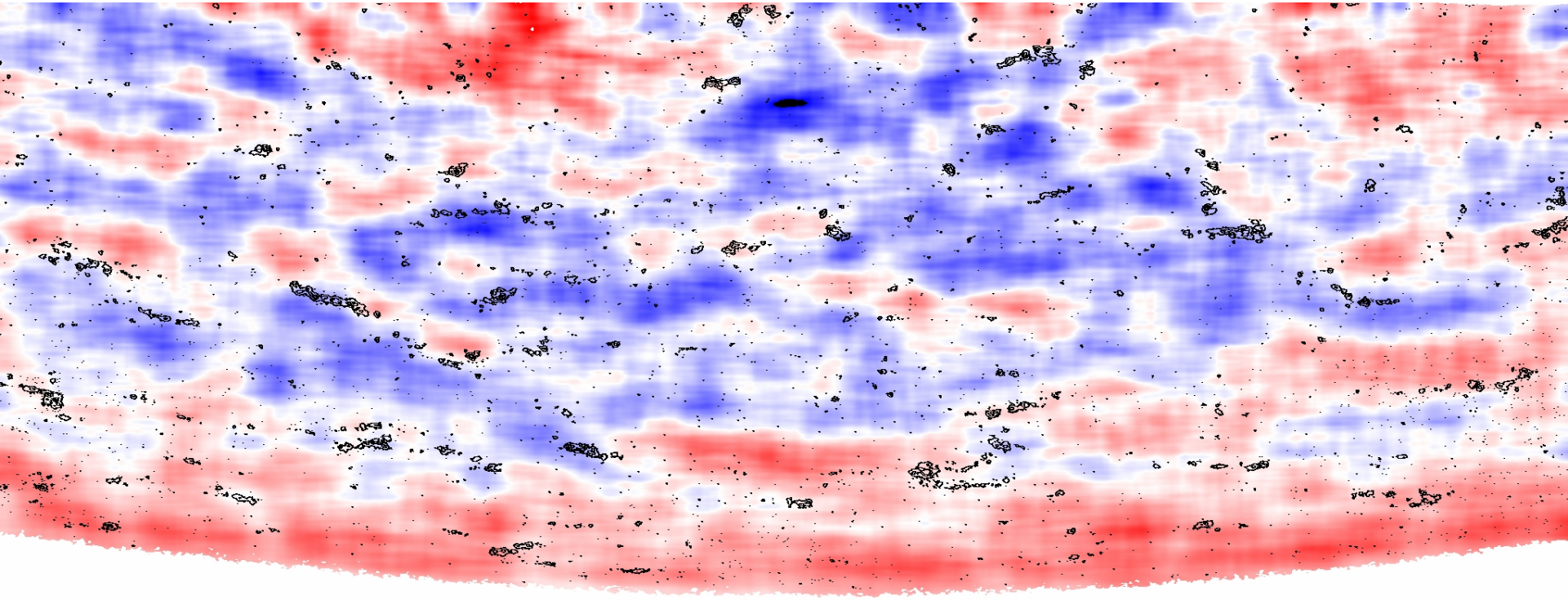
- Exact balance of north and south total flux
- Constant flux over 1 year

Temporal evolution of kG patches

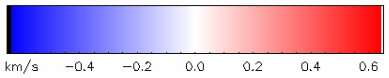
- kG patches have canopy structure
- Life time ~ 10-20 hours l
- Super Equi-partition & unipolar
- Small flux tubes merge to form large patch
- Large patch disintegrates to smaller patches



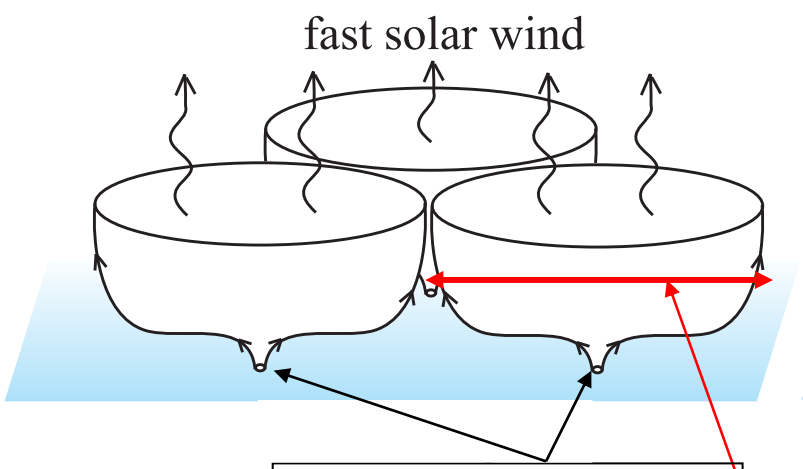
Velocity map obtained from Stokes-I and kG patches (contour)



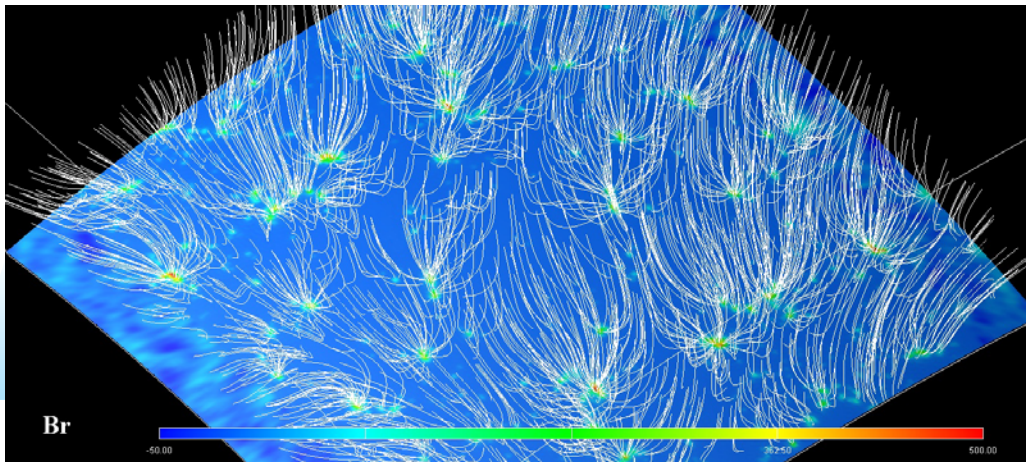
- **Relation between flow-field and magnetic patches are being analyzed.**



What is the implication of the kG polar field for acceleration of fast solar wind

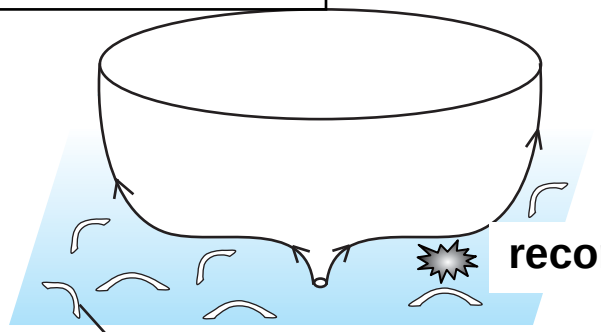


Fast solar wind comes from here



Polar field extrapolation (potential) using polar SP data embedded in MDI data (Shiota, Ito, Tsuneta, 2009)

Flares and jets



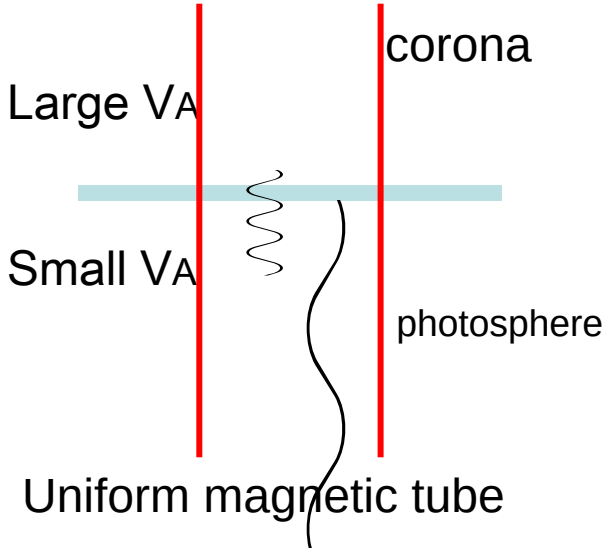
A factor of 345 inflation of the flux tubes near photosphere-choromosphere-corona boundary

Figure Courtesy Joten Okamoto

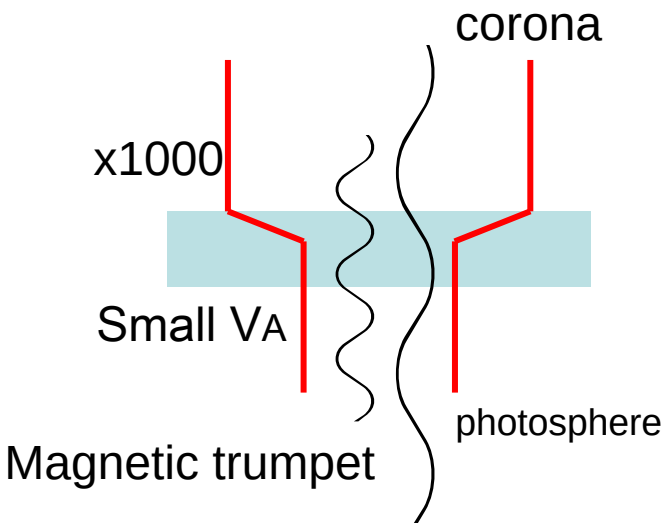
Fanning-out kG patch Chimney for Alfven waves

- ***Uniform magnetic field***: Aflven wave reflection in the photosphere-corona boundary

$$V_A = \frac{B}{\sqrt{4\pi\rho}}$$



- ***Fanning-out flux tube***: serves as chimney for Alfven waves
  - more Alfven wave flux that accelerates fast solar wind is transmitted to the corona due to the fanning-out structure.



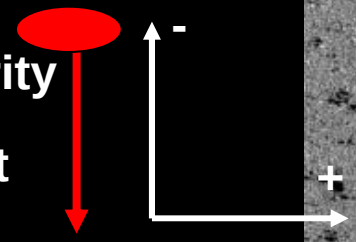
Summary
Properties of polar fields

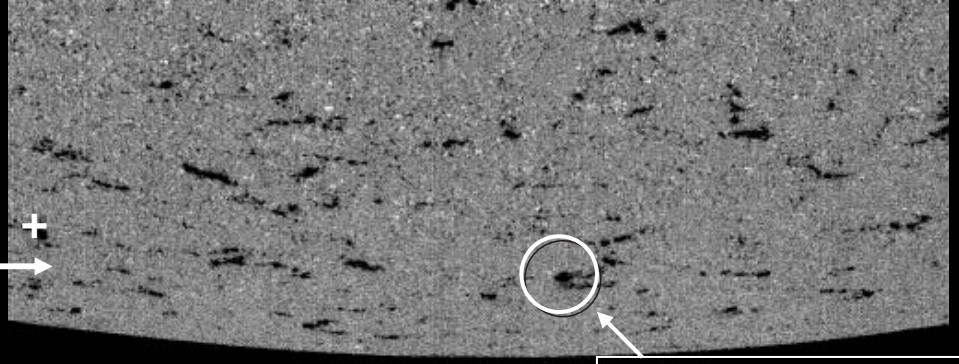
- *Vertical* field component
  - PDF with extenstion to super-equipartition (400-500G) field strength seen as kG patch
  - kG patch fanning-out structure
  - Unipolar in contrast to QS
  - Appear from nothing to disappear to nothing
  - 10-15 hour life time
  - North flux same as south flux
- Ubiquitous *horizontal* fields component everywhere as seen in Stokes V
  - PDF similar to those of quiet sun and plage region
  - Local dynamo process (Ishikawa and Tsuneta 2009)

Impact to solar dynamo

- Magnetic flux is transported to the polar regions with meridional flows and supergranular diffusion. Flux transport would be done via an aerodynamic (drag) force against magnetic tension force, and may be more difficult than the case for the mean field case.
- If the flux tubes seen on the surface of the Sun are maintained inside the Sun, they may affect a known difficulty in Ω-mechanism to generate intense toroidal field
  - an amplification factor of 100 that is needed to explain a combination of 1kG poloidal field with 100kG toroidal field may be achievable within a solar cycle.
  - There remains, however, a serious energetic problem.

Stokes Q black polarity
=North-South
magnetic component

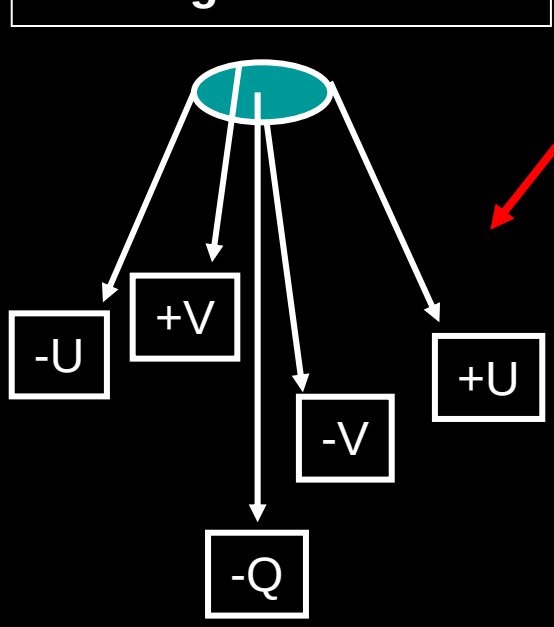


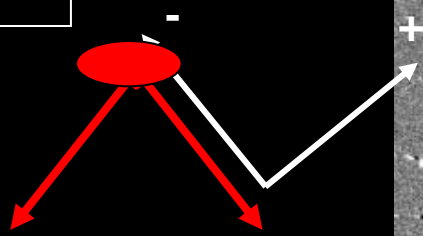


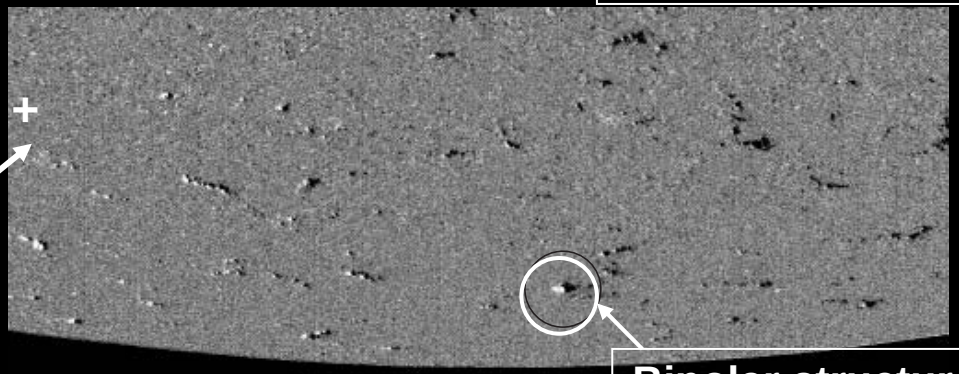
(b) Stokes Q

Unipolar structure
in Stokes Q

Polar flux tube
Fanning-out structure

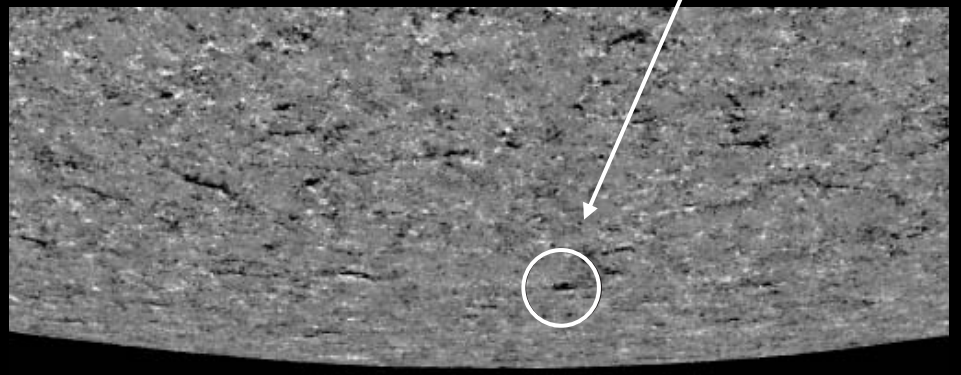






(c) Stokes U

Bipolar structure
in Stokes U & V

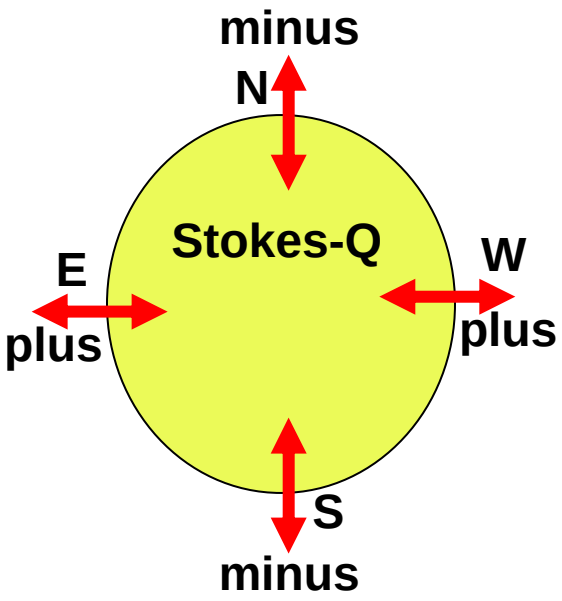


(d) Stokes V

Stokes V dominant black
polarity =magnetic field
away from the Sun

Milne-Eddington least-squares fit

- To minimize the influence of noise, fitting is performed for pixels whose *Q, U or V signals are larger than 5 sigma.*
  - 10.5% of the area meets the criteria.
- Formation layer at 80 degree latitude is approx. 100 km above the nominal height.
- *ME works fine for the extreme limb!*
- Parameters
  - vector magnetic field (strength, inclination angle, and azimuth angle), the line of sight velocity, two parameters describing the source function, the line-to-continuum absorption coefficient ratio, Doppler width, damping parameter, stray light factor(filling factor)


minus
N
Stokes-Q
E
plus
W
plus
S
minus